

Shaping the X-ray spectrum of galaxy clusters with AGN feedback and turbulence

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ABSTRACT

The hot plasma filling galaxy clusters emits copious X-ray radiation. The classic unheated and unperturbed cooling flow model predicts dramatic cooling rates and an isobaric X-ray spectrum with constant differential luminosity distribution. The observed cores of clusters (and groups) show instead a strong deficit of soft X-ray emission: $dL_x/dT \propto (T/T_{\text{hot}})^{\alpha=2\pm1}$. Using 3D hydrodynamic simulations, we show that such deficit arises from the tight self-regulation between thermal instability condensation and AGN outflow injection: condensing clouds boost the AGN outflows, which quench cooling as they thermalize through the core. The resultant average distribution slope is $\alpha \simeq 2$, oscillating within the observed $1 < \alpha < 3$. In the absence of thermal instability, the X-ray spectrum remains isothermal ($\alpha \gtrsim 8$), while unopposed cooling drives a too shallow slope, $\alpha < 1$. AGN outflows deposit their energy inside-out, releasing more heat in the inner cooler phase; radially distributed heating alone induces a declining spectrum, $1 < \alpha < 2$. Turbulence further steepens the spectrum and increases the scatter: the turbulent Mach number in the hot phase is subsonic, while it becomes transonic in the cooler phase, making perturbations to depart from the isobaric mode. Such increase in $d \ln P / d \ln T$ leads to $\alpha \approx 3$. Self-regulated AGN outflow feedback can address the soft X-ray problem through the interplay of heating and turbulence.

Key words: hydrodynamics – instabilities – turbulence – galaxies: active – galaxies: clusters: intracluster medium – galaxies: jets.

1 INTRODUCTION

Unbalanced radiative cooling of plasma would lead to a catastrophic cooling flow (CF) in the core of massive galaxies, groups, and clusters ($r < 0.1 R_{500}$). In galaxy clusters, the rate of the cooling gas would reach unrealistic values up to $\dot{M} \sim 10^3 M_\odot \text{ yr}^{-1}$ (Fabian 1994; Peres et al. 1998). The central cooling gas would monolithically condense out of the hot phase, leading to a dramatic increase in the core density and to a steady inflow due to the loss of pressure support against gravity. In the standard CF model, the CF is assumed to develop in isobaric mode. The mass deposition (or inflow) rate is thus independent of temperature (§3) and associated with a constant differential luminosity distribution

$$\frac{dL}{dT} = \frac{5}{2} \frac{\dot{M} k_b}{\mu m_p}, \quad (1)$$

with atomic weight $\mu \simeq 0.62$ for the intracluster medium (ICM).

In the last two decades, numerous observations in X-ray, UV, optical, and radio band have ruled out the presence of standard CFs in observed systems (Peterson & Fabian 2006; Böhringer & Werner 2010). The major CF issues are twofold: the mass sink problem and the soft X-ray problem. The former is the most studied problem, arising from the fact that we do not observe very large amount of warm ($H\alpha$, [NII]; McDonald et al. 2011; Werner et al. 2013) and cold molecular gas (CO , H_2 ; Edge 2001; Salomé & Combes 2003) dropped out of the hot phase. A

pure CF of almost $10^3 M_\odot \text{ yr}^{-1}$ would indeed double the mass of the central galaxy in less than 1 Gyr, quickly reaching $M_{\text{cold}} \sim 10^{12} M_\odot$ (Silk 1976; Cowie & Binney 1977; Mathews & Bregman 1978; Ciotti et al. 1991; Fabian 1994 for a comprehensive review on classic CF). Tightly related to the mass sink problem, the star formation rates appear to be quenched within a few per cent of the pure cooling rate (O’Dea et al. 2008).

The previous observations indicate that the pure cooling rates must be suppressed by one to two orders of magnitude. Early works in the 1990s (e.g. Binney & Tabor 1995; Ciotti & Ostriker 1997; Tucker & David 1997; Silk & Rees 1998) suggested that the required heating comes from the active galactic nucleus (AGN). With the advent of *XMM-Newton* and *Chandra*, it became clear that the AGN feedback – with powers up to $10^{46} \text{ erg s}^{-1}$ – plays the fundamental role in the thermodynamical evolution of gaseous atmospheres (McNamara & Nulsen 2007, 2012). In the last years, we tested different self-regulated AGN feedback models and compared them with data (Gaspari et al. 2009, 2011a,b, 2012b,a). The best consistent model is mechanical AGN feedback, in the form of massive subrelativistic outflows and self-regulated by the accretion of cold gas on to the black hole (Gaspari et al. 2013a, 2014b). The cooling rates are quenched by 20 fold, at the same time preserving the cool-core structure (as the positive temperature gradient). The rapid intermittency and high accretion rates of cold accretion permit the recurrent generation of buoyant bub-

bles, weak shocks, and the uplift of metal-rich low-entropy gas. The driven AGN turbulence is crucial for the growth of nonlinear thermal instability (TI): the condensation of extended filaments boosts the accretion rate, closing the self-regulated feedback loop (known as chaotic cold accretion mechanism – CCA; Gaspari et al. 2013b, 2014c). Such results have been independently corroborated by other observational and numerical investigations (Sharma et al. 2012; Li & Bryan 2014; Voit & Donahue 2015; Voit et al. 2015).

While the mass sink problem has been extensively tackled, the problem related to the deficit of emission in the soft X-ray band remains less studied. The high spectral resolution of *XMM* Reflection Grating Spectrometer (RGS) has revealed that the residual CF is not constant as a function of temperature. The emission measure of the X-ray spectrum shows a substantial deficit below 2 keV, typically correlated with the presence of multiphase structures (Kaastra et al. 2001; Peterson et al. 2001; Tamura et al. 2001; Sakelliou et al. 2002). Combining the observations of the last decade (see §3), the X-ray spectrum written as $dL_x/dT \propto (T/T_{\text{hot}})^\alpha$ is constrained to have a slope $\alpha \approx 2 \pm 1$ (Peterson et al. 2003; de Plaa et al. 2004; Kaastra et al. 2004; Sanders et al. 2009; Grange et al. 2010; Sanders et al. 2010).

In this work, we probe the impact of self-regulated AGN feedback in shaping the X-ray spectrum through the combined action of heating and turbulence. Direct heating reduces the average cooling rate. However, how can it deposit more energy in the cooler and denser gas with increasing emissivity? Does turbulence play a role in the cold phase thermodynamics? We propose that isobaric radiative cooling is inapplicable in a turbulent ICM atmosphere, even if motions are subsonic (§3). Both simulations and observations have shown that the ICM is turbulent (Schuecker et al. 2004; Nagai et al. 2007; Vazza et al. 2009; Gaspari & Churazov 2013; Gaspari et al. 2014a). Turbulence not only plays central role in the formation of nonlinear TI and dense filaments, but also makes the cooler phase to depart from the isobaric mode. As motions become transonic in the cold phase, this effect becomes increasingly more important at lower temperatures, contributing to the observed decline in the soft X-ray spectrum.

2 PHYSICS & NUMERICS

The setup and equations of the 3D hydrodynamic simulations with self-regulated AGN feedback are described in Gaspari et al. (2012b; hereafter G12). Here, we summarize the essential physics and numerics.

We adopt as reference system a massive galaxy cluster with virial mass $\approx 10^{15} M_\odot$ and baryon fraction 0.15. The system is modelled on the well-studied cluster A1795 by using the observed cool-core temperature profile (Ettori et al. 2002). The initial minimum core temperature is $\simeq 3$ keV; the self-regulated AGN feedback preserves it at ~ 3 –4 keV also during the evolution (the background profile related to the diffuse phase is thus not affecting the soft X-ray decline). The cluster is initialized in hydrostatic equilibrium in the Navarro-Frenk-White (NFW) potential dominated by dark matter. The initial profiles are perturbed (due to cosmological flows) to avoid idealized symmetric conditions, albeit the fluctuations are soon driven by the AGN turbulence. We use FLASH4 code with third-order accurate piecewise parabolic method (PPM) to integrate the hydrodynamic equations (equations 1–4 in G12). The domain is 1.27 Mpc on a side, covered by static mesh refinement zooming in the cluster core. The maximum resolution is 300 pc. The fuelling of the black hole is dominated by cold accretion (§1). Self-consistent accretion simulations down to 20 gravitational radii

showed that in CCA most of the cold gas condensing within 10 kpc is accreted in a few free-fall times (Gaspari et al. 2013b, 2014c), irrespectively of the Bondi radius, which is only meaningful for hot/Bondi accretion models. We ran a shorter simulation with double resolution and found very similar results, in particular in terms of the dL_x/dT diagnostic. The X-ray range ($\gtrsim 5 \times 10^6$ K) describing dL_x/dT (Fig. 1) is highly resolved. The smallest 10^4 K clouds approaching resolution may be diffused, but these have no impact on dL_x/dT and are negligible for the total accretion rate.

In addition to hydrodynamics, we include two key physical source terms: radiative cooling and self-regulated AGN outflows. The ICM emits radiation via bremsstrahlung ($T > 1$ keV) and line cooling ($T < 1$ keV), with total emissivity $\mathcal{L} = n_e n_i \Lambda(T, Z)$, where n_e and n_i are the electron and ion number density, respectively ($\mu_e \mu_i \simeq 1.464$). The cooling function $\Lambda(T, Z)$ is based on the tabulated values by Sutherland & Dopita (1993) for a fully ionized plasma with metallicity $Z = 0.3 Z_\odot$. The stable cold phase is preserved at 10^4 K. The split source term is integrated with an exact method, which tracks the cooling gas with high accuracy.

Unopposed radiative cooling would lead to a massive CF catastrophe. As shown by observations and simulations (§1), mechanical AGN feedback is the preferred mechanism to quench CFs (and subsequent star formation). Bipolar AGN outflows are injected through the internal boundaries in the middle of the domain with kinetic power

$$P_{\text{out}} = \epsilon \dot{M}_{\text{acc}} c^2, \quad (2)$$

where ϵ is the kinetic efficiency and \dot{M}_{acc} is the black hole accretion rate tied to the sinked gas in the inner zones. As shown in Gaspari et al. (2011a, 2014b), and G12, the long-lived presence of observed cool cores and the absence of severe overheating (negative $T(r)$ gradients) put tight constraints on the consistent kinetic efficiency, with best value $\epsilon \simeq 5 \times 10^{-3}$. Besides simulations (see also Di Matteo et al. 2005), observational works studying the AGN kinetic luminosity function also retrieve $\epsilon \simeq 5 \times 10^{-3}$ (cf. Merloni & Heinz 2008; La Franca et al. 2010). Most of the accreted gas is in the form of cold clouds and filaments, while the hot phase accounts for just a few per cent of the accretion rate (CCA; Gaspari et al. 2013b). The typical mass rate and velocity of the outflows are a few $M_\odot \text{ yr}^{-1}$ and $5 \times 10^4 \text{ km s}^{-1}$, respectively.

To disentangle the dominant physics, we perform two additional runs: one with purely distributed heating and one with distributed heating plus forced turbulence. Computationally, we set the heating rate (per unit volume) to be equal to the average radiative emissivity in finite radial shells, $\mathcal{H} \approx \langle \mathcal{L} \rangle$. Forced turbulence is driven via a spectral forcing scheme based on an Ornstein-Uhlenbeck random process, which drives a time-correlated and zero-mean acceleration field. The driven velocity dispersion is $\sigma_v \sim 250 \text{ km s}^{-1}$ (with injection scale $L \approx 80 \text{ kpc}$). Both modules are described in detail in Gaspari et al. 2013b (section 2).

3 RESULTS & DISCUSSION

In Fig. 1, we present the key diagnostic plot defining the X-ray spectrum: the differential luminosity distribution per temperature interval. The dL_x/dT distribution contains crucial information. The energy equation (or first law of thermodynamics) affected by radiative losses and heating can be written in Lagrangian form as

$$\rho \frac{d\varepsilon}{dt} - \frac{P}{\rho} \frac{d\rho}{dt} = -n_e n_i \Lambda + n \Gamma, \quad (3)$$

where $\varepsilon = k_b T / [\mu m_p (\gamma - 1)]$ is the gas specific thermal energy, $P = k_b \rho T / \mu m_p$ is the pressure, $\gamma = 5/3$ is the adiabatic index,

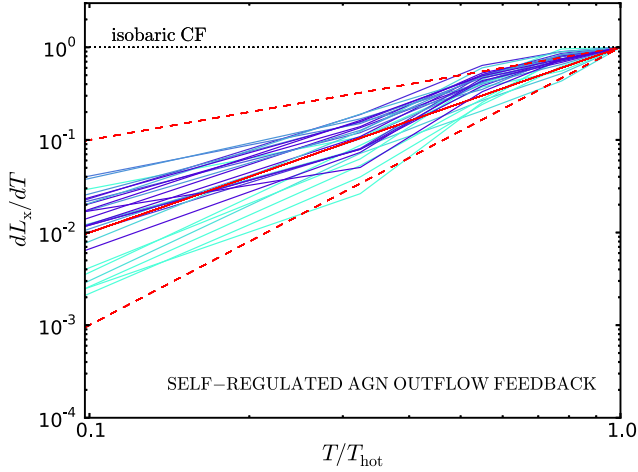


Figure 1. The X-ray spectrum in the form of differential luminosity distribution per temperature interval, for the reference self-regulated AGN outflow feedback model ($\epsilon \simeq 5 \times 10^{-3}$). As in observations (§1), we use the hot gas temperature in the core for the normalization ($T_{\text{hot}} \simeq 4$ keV) and we extract the distribution within the cooling radius, $r < 100$ kpc (~ 1 arcmin for A1795). The distribution is normalized to the pure isobaric CF (dotted; equation 6). The colour coding is related to the accretion rates averaged over 20 Myr after each snapshot time (cyan to blue ≈ 1 –20 $M_{\odot} \text{ yr}^{-1}$). The simulated spectrum is consistent with the observational constraints (red; $\alpha = 2 \pm 1$, equation 9), showing a strong deficit down to $T \simeq 0.4$ keV (we cut the plot to this value as it roughly corresponds to *Chandra* lower sensitivity). The departure from the standard isobaric CF naturally develops via the tight self-regulation between TI condensation and turbulent AGN heating.

and $\Gamma \equiv dE/dt$ is the injected heating rate per particle (erg s^{-1}). We focus in this study on the differential X-ray luminosity, which is defined as $dL_x \equiv n_e n_i \Lambda(T, Z) dV$ (Peterson & Fabian 2006), where $\Lambda(T, Z)$ is the cooling function (see §2). Using the classic definition $\dot{M} \equiv -\rho dV/dt$ in equation (3), yields

$$\frac{dL_x}{dT} = \dot{M} \left(\frac{\gamma}{\gamma - 1} \frac{k_b}{\mu m_p} - \frac{1}{\rho} \frac{dP}{dT} - \frac{1}{\mu m_p} \frac{dE}{dT} \right) \quad (4)$$

$$= \frac{\dot{M} k_b}{\mu m_p} \left(\frac{5}{2} - \gamma_T - \frac{1}{k_b} \frac{dE}{dT} \right), \quad (5)$$

where $\gamma_T \equiv d \ln P / d \ln T = \gamma_{\text{eff}} / (\gamma_{\text{eff}} - 1)$; $\gamma_{\text{eff}} = d \ln P / d \ln \rho$ is known as the effective adiabatic index. Equation (5) indicates dL_x/dT can be reduced by changing the effective adiabatic index (e.g. turbulence) and/or an external injection of heat (last term).

Let us neglect the heating term for the moment ($dE = 0$). If cooling proceeds in isobaric mode (as in the standard CF), then $\gamma_T = 0$ ($\gamma_{\text{eff}} = 0$) and the luminosity distribution in equation (5) depends only on a given mass inflow rate (dotted horizontal line in Fig. 1):

$$\frac{dL_x}{dT} = \frac{5}{2} \frac{\dot{M} k_b}{\mu m_p} \quad [\text{isobaric}]. \quad (6)$$

Note that the X-ray spectrum can be interpreted in terms of differential emission measure $dY/dT = \Lambda^{-1} dL_x/dT$. In the isobaric CF, $dY/dT \propto T$, since $\Lambda \propto T^{-1}$ and $dL/dT \propto \dot{M}$ is constant. If cooling develops instead in isochoric mode (constant density), $\gamma_T = 1$ ($|\gamma_{\text{eff}}| = \infty$), the luminosity distribution is reduced to

$$\frac{dL_x}{dT} = \frac{3}{2} \frac{\dot{M} k_b}{\mu m_p} \quad [\text{isochoric}]. \quad (7)$$

In the extreme adiabatic mode, $\gamma_T = 5/2$ ($\gamma_{\text{eff}} = 5/3$), the luminosity distribution can be strongly suppressed:

$$\frac{dL_x}{dT} \ll \frac{\dot{M} k_b}{\mu m_p} \quad [\text{adiabatic}]. \quad (8)$$

Which are the constraints provided by observations? Kaastra et al. (2004) studied dY/dT in the core of 17 well-observed CF clusters. The average sample slope of the emission measure distribution is $\alpha' \approx 3.1$, which translates to a luminosity slope $\alpha = \alpha' - 1 \approx 2.1$. The scatter is significant, with most clusters contained within $1 < \alpha < 3$. A few objects have extremely low and high values, such as Perseus ($\alpha \approx 0.5$), A1835 ($\alpha \sim 0$), MKW9 ($\alpha \approx 8$), and A399 ($\alpha \rightarrow \infty$, isothermal). Peterson et al. (2003) also rule out the standard CF, finding multitemperature spectra with $\alpha \approx 1$ –2 in 14 clusters (fig. 9). de Plaa et al. (2004) measure $\alpha \approx 1.63$ in the core of A478. Grange et al. (2010) retrieve $\alpha \approx 1.1$ in J1539 cluster. Sanders et al. (2009) find a soft X-ray deficit consistent with $\alpha \sim 1.5$ in 2A 0335+096. The X-ray bright cool-core clusters A262, A3581, and HCG62 show a multiphase structure with $\alpha \sim 2$ (Sanders et al. 2010), although the distribution is more complex than a simple power law. In Centaurus cluster, the RGS spectrum has $\alpha \approx 1.31$ (Sanders et al. 2008), while M87 spectrum is fitted by $\alpha \approx 1.13$ (Werner et al. 2006). Massive groups similarly display suppressed soft X-ray emission. For instance, in NGC 5044, Tamura et al. (2003) retrieve $\alpha \sim 3$ below 1 keV, flattening near 0.3 keV. To summarize, the observed deficit in the soft X-ray luminosity distribution can be written as

$$\frac{dL_x}{dT} \simeq \frac{5}{2} \frac{\dot{M} k_b}{\mu m_p} \left(\frac{T}{T_{\text{hot}}} \right)^{\alpha}, \quad (9)$$

where $\alpha = 2 \pm 1$ (red lines in Figs 1 and 2) and T_{hot} is the ambient hot temperature in each cluster core.

In Fig. 1 (blue and cyan), we show the differential X-ray luminosity distribution resulting from the reference AGN feedback simulation during tight self-regulation (§2), extracted in the cluster core ($r < 100$ kpc) and sampled every 50 Myr. The differential luminosity decreases from the core ambient temperature ($T_{\text{hot}} \simeq 4$ keV) to the soft X-ray floor ($\simeq 0.4$ keV), reaching a deficit of 2 dex compared with the pure isobaric CF (dotted black). The simulated dL_x/dT slope oscillates around $\alpha \simeq 2$, with a scatter contained within 1 and 3, which is consistent with the observed values (equation 9). The steady oscillation around $\alpha \approx 2$ is reached when the system is tightly self-regulated, dancing between periods of TI condensation and outflow injection. Perturbations in the heated atmosphere grow quickly nonlinear as the ‘TI-ratio’ $t_{\text{cool}}/t_{\text{ff}} < 10$ (G12, sec. 2), i.e. in the region $r \lesssim 20$ kpc for A1795. Cold clouds and extended filaments condense out of the hot phase, rain on to the black hole, boosting the accretion rate and the AGN outflow power. Correlating dL_x/dT with the accretion rate averaged over 20 Myr after each snapshot time, we find that higher values are linked to shallower dL_x/dT . Since the evolution is self-regulated, we are witnessing the continuous fuelling of the black hole via cold accretion. The stronger the TI condensation (lower α), the larger the amount of fuel, subsequently boosting P_{out} (up to $\sim 10^{45} \text{ erg s}^{-1}$).

We analysed other runs departing from proper self-regulation. In the absence of TI (e.g. in runs with $\epsilon > 0.01$ or $t_{\text{cool}}/t_{\text{ff}} \gg 20$), the X-ray spectrum approaches the isothermal behaviour with $\alpha \approx 8$ (not shown). In the other extreme run with pure cooling, we find a shallow spectrum slope, $\alpha \approx 0.7$. We notice that even without heating, a cooling halo develops nonlinear TI in the presence of relevant perturbations (e.g. if the amplitude is > 0.1 ; Gaspari et al. 2013b, 2014c). The TI condensation process and the preexisting

perturbations drive chaotic motions ($\sigma_v \sim 70 \text{ km s}^{-1}$) leading to the departure from the pure isobaric mode ($\alpha > 0$; see below). Moreover, below 1 keV, t_{cool} starts to fall below the sound crossing time t_{cs} , helping the perturbation mode to shift to the isochoric regime (equation 7). A soft X-ray decline is thus associated with the formation of TI and a multiphase cluster core.

The irreversible heating due to the AGN feedback is a key component suppressing the luminosity distribution (last term in equation 5). Why should dE/dT increase at lower T ? Simulations show that AGN feedback is an inside-out process, depositing more energy in the inner regions and less energy at larger radii. The AGN outflows are injected through the inner boundary and the simulations self-consistently compute the complex nonlinear thermalization of the outflows through the cluster core. As shown in G12 (fig. 9, right), such energy deposition is maximal in the central few kpc and decreases with radial distance (bottom to top panels). Below, we test forced radially distributed heating, which results in a similar dL_x/dT , further corroborating that the more realistic AGN outflow feedback simulations approach the inside-out heating distribution. Since cooling strongly increases at smaller radii and condensed TI sink to the centre, AGN heating naturally targets the inner cooler regions, inducing larger dE/dT at lower temperatures, thereby steepening the differential luminosity distribution (Fig. 1).

On top of external AGN heating, is turbulence contributing to suppressing the soft X-ray spectrum? In the $dE = 0$ case, in order to match observations (equation 9), the γ_T index (equation 5) has to vary as

$$\gamma_T \approx \frac{5}{2} \left[1 - \left(\frac{T}{T_{\text{hot}}} \right)^\alpha \right], \quad (10)$$

i.e. γ_T should monotonically increase with decreasing temperature. In other words, the gas dynamics should shift from the isobaric towards the isochoric/adiabatic mode at low T . It has been shown that the thermodynamic mode driven by turbulence changes with the Mach number (Gaspari & Churazov 2013; Gaspari et al. 2014a), shifting from the isobaric to adiabatic mode as $\text{Ma} > 0.5$. The subsonic motions driven in the hot phase become transonic/supersonic in the condensing cooler phase. This increases the importance of the non-isobaric modes (see also interstellar medium studies; Sánchez-Salcedo et al. 2002), together with the impact of turbulent mixing, which can further steepen dL_x/dT .

Since the self-regulated AGN outflow models include a complex nonlinear mix of heating and turbulence, we computed two additional control runs: one purely with heating distributed in radial shells (now an assumption, not a result) and the other with the addition of forced subsonic turbulence ($\sigma_v \sim 250 \text{ km s}^{-1}$; §2). Fig. 2 shows that direct heating (dark blue lines) is important to induce a decline of the spectrum with $1 < \alpha < 2$, thereby reproducing part of the AGN feedback evolution (Fig. 1). However, turbulence (orange) is necessary to broaden the distribution (up to 1 dex at low T) and to generate steeper X-ray spectra with slopes up to $\alpha \approx 3$. The increase of γ_T at lower temperature (equation 10) via turbulence is thus an important component to address the soft X-ray problem.

Returning to the reference self-regulated run (Fig. 1), we computed the vorticity magnitude in the cluster core for two random snapshots having $\alpha \approx 1.5$ and 2.5. The only source of vorticity is the turbulence excited by the AGN feedback (via outflows, bubbles, and shocks); thereby, the driven vorticity $\omega \equiv \nabla \times \mathbf{v}$ is a good tracer of turbulent motions. From the shallow to steeper curve, we find that the vorticity magnitude increases from $\omega \approx 0.02$ to 0.04 Myr^{-1} . The latter value is analogous to the vorticity retrieved in the forced turbulence run with $\sigma_v \sim 250 \text{ km s}^{-1}$. Therefore, a

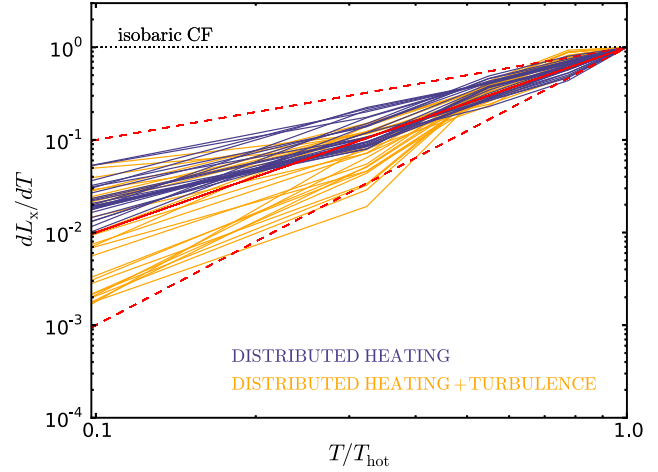


Figure 2. Differential X-ray luminosity distribution for the run with distributed heating (no turbulence; dark blue) and the same model with the addition of forced turbulence ($\sigma_v \sim 250 \text{ km s}^{-1}$; orange). Analogue of Fig. 1. The average cooling rate is comparable to the outflow feedback models, $\approx 30 \text{ M}_\odot \text{ yr}^{-1}$. Direct heating is fundamental to produce a substantial decline in the soft X-ray spectrum. Since the AGN heating distribution is a decreasing function of radius, it is naturally tied to the cooler, inner gas, leading to a positive slope α . Turbulence increases the importance of adiabatic processes and mixing in the cooler phase (where it becomes transonic), further widening the departure from the isobaric regime ($\gamma_T > 0$; equation 5). The γ_T increase steepens on average the spectrum slope, with values up to $\alpha \approx 3$, while the scatter is now raised to ~ 1 dex at low temperatures.

subsonic turbulent Mach number in the hot phase (4 keV) as small as $\text{Ma}_{\text{hot}} \sim 0.25$ can have important consequences for the condensing cooler phase ($\text{Ma}_{\text{cold}} \gtrsim 1$), leading to departures from the isobaric mode and a constant differential luminosity, even in the pure cooling runs. As further support for this scenario, we find in all runs a systematic drop of pressure in the cold phase down to 10^4 K . While the pure adiabatic mode ($\gamma_T = 5/2$) is hard to achieve, the isochoric regime is the most frequent state of the cold phase. Turbulence thus requires the complementary AGN heating to induce substantially steep X-ray spectra (Fig. 2).

4 CONCLUSIONS

This work aims to study the impact of AGN feedback and turbulence in shaping the X-ray spectrum of galaxy cluster cores, in particular addressing the observed deficit of emission towards the soft X-ray band. We combined extensive literature data of the last decade, showing that the observed X-ray differential luminosity distribution has a steep decline constrained as $dL_x/dT \propto (T/T_{\text{hot}})^{\alpha=2\pm 1}$. This is a dramatic deviation from the standard isobaric CF model having constant distribution (equation 6). The 3D hydrodynamic simulation with AGN outflow feedback ($\epsilon \simeq 5 \times 10^{-3}$) reproduces the observed soft X-ray spectrum decline during tight self-regulation, $1 < \alpha < 3$. The simulations show that the deficit naturally arises from the symbiotic link between cold gas condensation and AGN outflow injection: in a turbulent and heated atmosphere, thermal instabilities grow quickly nonlinear, boosting the accretion rate and thus the AGN outflows; the kinetic outflows gradually thermalize through the core, reheating the gas and thus quenching cooling, in a self-regulated loop.

Without tight self-regulation, two scenarios are ruled out by soft X-ray observations. If the system is strongly overheated (e.g. due to very high feedback efficiency or $t_{\text{cool}}/t_{\text{ff}}$), cooling and

growing TI are negligible, thereby the spectrum remains nearly isothermal ($\alpha \gtrsim 8$). Conversely, with negligible AGN heating, a pure CF develops producing a shallow spectrum slope, $\alpha < 1$. Even in the latter pure cooling regime, turbulent motions excited by the TI condensation and cosmological perturbations prevent to reach the isobaric mode. The drop in the soft X-ray band is thus associated with the formation of TI and multiphase cores.

To disentangle the main processes, we performed two additional control runs forcing radially distributed heating and subsonic turbulence, which are self-consistently achieved in the more realistic AGN outflow runs. Radially distributed heating alone produces a declining X-ray spectrum with slope $1 < \alpha < 2$. The heat deposition is higher for lower T gas, since the AGN outflows thermalize as a decreasing function of radius (G12), in similar way as the forced distributed heating models. Therefore, the AGN energy is deposited relatively more in the inner cooler phase (TI condensation further advects the cold gas to the centre), producing a steep spectrum. Forced subsonic turbulence ($\sigma_v \sim 250 \text{ km s}^{-1}$) introduces a significant scatter (1 dex) and further steepens the spectrum to $2 < \alpha < 3$. While the subsonic Mach number in the hot phase becomes transonic in the cooler phase, the perturbation mode departs from isobaric ($\gamma_T > 0$), suppressing dL_\times/dT at progressively lower T (equation 5). The self-regulated feedback simulation similarly shows that the slope steepens with increasing vorticity (tracer of turbulence), with maximum suppression reached with the vorticity of the forced stirring run. This correlation and the drop of pressure in the cold phase corroborate the scenario of non-isobaric cooling.

To conclude, mechanical AGN feedback tightly self-regulated via CCA is able to quench the cooling rates by 20 fold while preserving the cool-core structure (the mass sink problem); at the same time, such mechanism can induce a steep X-ray spectrum (the soft X-ray problem) through the complementary action of AGN outflow heating and the driven subsonic turbulence.

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REFERENCES

- Binney J., Tabor G., 1995, *MNRAS*, 276, 663
 Böhringer H., Werner N., 2010, *A&AR*, 18, 127
 Ciotti L., D’Ercole A., Pellegrini S., Renzini A., 1991, *ApJ*, 376, 380
 Ciotti L., Ostriker J. P., 1997, *ApJ*, 487, L105
 Cowie L. L., Binney J., 1977, *ApJ*, 215, 723
 de Plaa J., Kaastra J. S., Tamura T., Pointecouteau E., Mendez M., Peterson J. R., 2004, *A&A*, 423, 49
 Di Matteo T., Springel V., Hernquist L., 2005, *Nature*, 433, 604
 Edge A. C., 2001, *MNRAS*, 328, 762
 Ettori S., Fabian A. C., Allen S. W., Johnstone R. M., 2002, *MNRAS*, 331, 635
 Fabian A. C., 1994, *ARA&A*, 32, 277
 Gaspari M., Churazov E., 2013, *A&A*, 559, A78
 Gaspari M., Melioli C., Brighenti F., D’Ercole A., 2009, in S. Heinz, E. Wilcots, eds, *AIP Conf. Proc. Vol. 1201, The Monster’s Fiery Breath: Feedback in Galaxies*. Am. Inst. Phys., New York, p. 309
 Gaspari M., Melioli C., Brighenti F., D’Ercole A., 2011a, *MNRAS*, 411, 349
 Gaspari M., Brighenti F., D’Ercole A., Melioli C., 2011b, *MNRAS*, 415, 1549
 Gaspari M., Brighenti F., Temi P., 2012a, *MNRAS*, 424, 190
 Gaspari M., Ruszkowski M., Sharma P., 2012b, *ApJ*, 746, 94
 Gaspari M., Brighenti F., Ruszkowski M., 2013a, *AN*, 334, 394
 Gaspari M., Ruszkowski M., Oh S. P., 2013b, *MNRAS*, 432, 3401
 Gaspari M., Churazov E., Nagai D., Lau E. T., Zhuravleva I., 2014a, *A&A*, 569, A67
 Gaspari M., Brighenti F., Temi P., Ettori S., 2014b, *ApJ*, 783, L10
 Gaspari M., Ruszkowski M., Oh S. P., Brighenti F., Temi P., 2014c, preprint (arXiv:1407.7531)
 Grange Y. G., Costantini E., de Plaa J., in’t Zand J. J. M., Verbunt F., Kaastra J. S., Verrecchia F., 2010, *A&A*, 513, A2
 Kaastra J. S., Ferrigno C., Tamura T., Paerels F. B. S., Peterson J. R., Mittaz J. P. D., 2001, *A&A*, 365, L99
 Kaastra J. S. et al., 2004, *A&A*, 413, 415
 La Franca F., Melini G., Fiore F., 2010, *ApJ*, 718, 368
 Li Y., Bryan G. L., 2014, *ApJ*, 789, 54
 McDonald M., Veilleux S., Mushotzky R., 2011, *ApJ*, 731, 33
 McNamara B. R., Nulsen P. E. J., 2007, *ARA&A*, 45, 117
 McNamara B. R., Nulsen P. E. J., 2012, *New J. Phys.*, 14, 055023
 Mathews W. G., Bregman J. N., 1978, *ApJ*, 224, 308
 Merloni A., Heinz S., 2008, *MNRAS*, 388, 1011
 Nagai D., Vikhlinin A., Kravtsov A. V., 2007, *ApJ*, 655, 98
 O’Dea C. P. et al., 2008, *ApJ*, 681, 1035
 Peres C. B., Fabian A. C., Edge A. C., Allen S. W., Johnstone R. M., White D. A., 1998, *MNRAS*, 298, 416
 Peterson J. R., Fabian A. C., 2006, *Phys. Rep.*, 427, 1
 Peterson J. R., Kahn S. M., Paerels F. B. S., Kaastra J. S., Tamura T., Bleeker J. M., Ferrigno C., Jernigan J., 2003, *ApJ*, 590, 207
 Peterson J. R. et al., 2001, *A&A*, 365, L104
 Sakelliou I. et al., 2002, *A&A*, 391, 903
 Salomé P., Combes F., 2003, *A&A*, 412, 657
 Sánchez-Salcedo F. J., Vázquez-Semadeni E., Gazol A., 2002, *ApJ*, 577, 768
 Sanders J. S., Fabian A. C., Allen S. W., Morris R. G., Graham J., Johnstone R. M., 2008, *MNRAS*, 385, 1186
 Sanders J. S., Fabian A. C., Frank K. A., Peterson J. R., Russell H. R., 2010, *MNRAS*, 402, 127
 Sanders J. S., Fabian A. C., Taylor G., 2009, *MNRAS*, 396, 1449
 Schuecker P., Finoguenov A., Miniati F., Böhringer H., Briel U. G., 2004, *A&A*, 426, 387
 Sharma P., McCourt M., Quataert E., Parrish I. J., 2012, *MNRAS*, 420, 3174
 Silk J., 1976, *ApJ*, 208, 646
 Silk J., Rees M. J., 1998, *A&A*, 331, L1
 Sutherland R. S., Dopita M. A., 1993, *ApJS*, 88, 253
 Tamura T., Kaastra J. S., Makishima K., Takahashi I., 2003, *A&A*, 399, 497
 Tamura T. et al., 2001, *A&A*, 365, L87
 Tucker W., David L. P., 1997, *ApJ*, 484, 602
 Vazza F., Brunetti G., Kritsuk A., Wagner R., Gheller C., Norman M., 2009, *A&A*, 504, 33
 Voit G. M., Donahue M., 2015, *ApJ*, 799, L1
 Voit G. M., Donahue M., Bryan G. L., McDonald M., 2015, *Nature*, 519, 203
 Werner N., Böhringer H., Kaastra J. S., de Plaa J., Simionescu A., Vink J., 2006, *A&A*, 459, 353
 Werner N. et al., 2013, *ApJ*, 767, 153